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14. ABSTRACT We document the detailed spatio-temporal structure of the dayside aurora during intervals of ongoing dayside magnetopause reconnection, primarily during interplanetary magnetic field (IMF) $B_z \lesssim 0$ conditions. The present study is based on ground auroral observations in combination with particle precipitation data from a DMSP spacecraft. We describe auroral forms corresponding to the following particle precipitation regimes identified by Newell and Meng (1994): (i) central plasma sheet (CPS), (ii) precipitation void, (iii) dayside boundary plasma sheet (BPS), and (iv) cusp (LLBL/cusp/mantle). Two distinctly different auroral configurations are observed, corresponding to different regimes of the IMF clock angle (θ) and the $|B_y|/|B_z|$ ratio. Two regimes are defined in regime (I) θ lies within ~ 90 - 135 deg and $|B_y|/|B_z| < 1$ (B-dominated), while in regime (II) θ is in the range 135 - 180 deg and $|B_y|/|B_z| > 1$ (B-dominated). Within regime (I) the auroral response to reconnection events typically progresses from lower to higher latitudes in stages as indicated below: (A) equatorward boundary intensifications (EBIs): sequential brightenings of closely spaced, fragmented, rayed bands (BPS aurora) within the 08:00-15:00 MLT sector, each of which are moving noonward/sunward, (B) poleward moving auroral forms (PMAFs): forms expanding westward from the expanding from in the convention throat in the 09:00-12:00 MLT sector, with a final phase in the region of mantle precipitation. During strongly southward IMF conditions (regime II) the intense PMAF activity is replaced by a more latitudinally restructured, but longitudinally wide auroral of moderate intensity. The latter auroral state is accompanied by a 2-cell convection pattern which is rather symmetrical about noon.

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Dayside aurora and the role of IMF $|B_y|/|B_z|$: detailed morphology and response to magnetopause reconnection

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Abstract. We document the detailed spatio-temporal structure of the dayside aurora during intervals of ongoing dayside magnetopause reconnection, primarily during interplanetary magnetic field (IMF) $B_z \leq 0$ conditions. The present study is based on ground auroral observations in combination with particle precipitation data from a DMSP spacecraft. We describe auroral forms corresponding to the following particle precipitation regimes identified by Newell and Meng (1994): (i) central plasma sheet (CPS), (ii) precipitation void, (iii) dayside boundary plasma sheet (BPS), and (iv) cusp (LLBL/cusp/mantle). Two distinctly different auroral configurations are observed, corresponding to different regimes of the IMF clock angle (θ) and the $|B_y|/|B_z|$ ratio. Two regimes are defined. In regime (I) θ lies within $\sim 90^\circ$ – 135° and $|B_y|/|B_z| > 1$ (B_y -dominated), while in regime (II) θ is in the range 135° – 180° and $|B_y|/|B_z| < 1$ (B_z -dominated). Within regime (I) the auroral response to reconnection events typically progresses from lower to higher latitudes in stages as indicated below: (A) equatorward boundary intensifications (EBIs): sequential brightenings of closely spaced, fragmented, rayed bands (BPS aurora) within the $\sim 08:00$ – $15:00$ MLT sector, each of which are moving noonward/sunward, (B) poleward moving auroral forms (PMAFs): forms expanding westward from the postnoon side ($B_y > 0$) and later appearing as a poleward expanding form in the convection throat in the $\sim 09:00$ – $12:00$ MLT sector, with a fading phase in the regime of mantle precipitation. During strongly southward IMF conditions (regime II), the intense PMAF activity is replaced by a more latitudinally restricted, but longitudinally wide aurora of moderate intensity. The latter auroral state is accompanied by a 2-cell convection pattern which is rather symmetrical about noon. This state is very different from the convection/FAC configuration present during IMF regime (I), with its strong zonal flows (convection current), more intense FAC sheets and PMAF activity in the midday sector. The strong

IMF regulation of the dayside BPS aurora, consisting of keV electrons, and its location with respect to the green line auroral gap (precipitation void), indicate that it is an important signature of the reconnection process, located on open boundary layer field lines. The observed longitudinal bifurcation of the auroral brightenings (EBIs) preceding PMAFs is consistent with antiparallel magnetopause reconnection.

Key words. Magnetospheric physics (auroral phenomena magnetopause, cusp, and boundary layers: solar wind-magnetosphere interactions) – Ionosphere (particle precipitation)

1 Introduction

Large statistical surveys have demonstrated that the various plasma populations of the dayside magnetosphere are manifested in corresponding particle precipitation regimes, as observed from satellites in polar orbit (Newell and Meng, 1988; Newell et al., 1989; Newell and Meng, 1994). While statistical satellite data gives the general (averaged) precipitation pattern, case studies of ground observations of the aurora provide detailed information on the spatio-temporal structure of the precipitation pattern within the restricted fields of view of the ground instruments. Specifically, the magnetospheric response to the variable solar wind/IMF conditions can be monitored from the local ground auroral observations.

The control of dayside auroral activity by IMF parameters, such as the clock angle (θ) and the IMF B_y polarity, has been shown in previous work (Sandholt et al., 1998b). The clock angle is the polar angle of the IMF in GSM Y – Z plane. Two latitudinally separate auroral forms, called type 1 and type 2, have been related to reconnection at sub-cusp latitudes (type 1), and poleward of the cusp (the higher latitude type 2 cusp aurora), respectively. The east-west motion of poleward moving auroral forms (PMAFs) is governed by IMF B_y and is toward dawn (dusk) when $B_y > 0$ (< 0), in the Northern Hemisphere (Sandholt et al., 1986, 1993). PMAFs

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are typically preceded by so-called equatorward boundary intensifications (EBIs) (Sandholt and Farrugia, 2003; Sandholt et al., 2003). In this study the detailed morphology and IMF regulation of the dayside aurora is probed further. We selected a 2-h long IMF interval on 8 December 1999, where (i) the IMF clock angle varies over its whole range (0° – 180°), (ii) both abrupt as well as smooth IMF changes occur, (iii) intervals of strongly northward and strongly southward IMFs are present, and (iv) there is an interval where the IMF lies in an “intermediate” (θ between 90° and 135°) state of orientation.

Using meridian scanning photometers (MSP) and all-sky camera (ASC) images we detail the response phenomenology in the dayside aurora, illustrating radically different reactions to these various IMF states/rotations. We achieve a comprehensive, detailed description of the auroral response to the IMF orientation, as parameterized by the clock angle (θ) and the ratio $|B_y/B_z|$, in the MLT region 09:00–12:00 MLT.

Particular attention is directed to the different auroral responses observed when the IMF points strongly south ($135^\circ < \theta < 180^\circ$) compared to that which we document for the “intermediate” clock angle regime ($90^\circ < \theta < 135^\circ$). We find that, in contrast to the intense and active response evident over a wide latitude range for the “intermediate” regime, the former state is characterized by the absence of PMAF transients. The aurora in this state (strongly south IMF) configures as a longitudinally wide band (expanded cusp) of moderate intensity, located, by comparison, at lower latitudes.

The dayside auroral morphology reported here involves the following particle precipitation regimes: (i) central plasma sheet (CPS), (ii) precipitation void, (iii) dayside boundary plasma sheet (BPS), and (iv) LLBL/cusp/mantle (Newell and Meng, 1994). The particle precipitation counterpart to the auroral forms is documented by combining ground and satellite observations, as has been done in previous studies (see, e.g. Denig et al., 1993).

Of great interest to us in this study is the role of the dayside BPS aurora in a reconnection setting. At this point we note that our dayside BPS corresponds to the particle precipitation regime referred to as the “soft electron zone” (SEZ) by Lyons et al. (1996). In previous studies it has been pointed out that events of transient magnetopause reconnection may leave an important imprint in the dayside BPS precipitation and its associated aurora (Sandholt and Newell, 1992; Sandholt et al., 1993; Kan et al., 1996; Oksavik et al., 2000). In this paper we document the strong and rapid IMF regulation of the dayside BPS aurora. This IMF regulation, in combination with the presence of a precipitation void (and associated gap in the green line auroral emission) at its equatorward edge, indicates that it is located on open field lines. The precipitation void between the CPS and BPS regimes is likely a signature of the escape of CPS electrons to the magnetosheath, along newly-opened field lines (Scholer et al., 1981, 1982). Our inference concerning the location of the BPS aurora with respect to the open-closed field line boundary is in agreement

with the reasoning of Lockwood et al. (1996) and Lockwood (1997).

According to the observations reported in this paper, the most spectacular auroral response to reconnection events occurs for IMF clock angles in the intermediate regime ($|B_y|/|B_z| \geq 1$). It consists of the following scenario in the $\sim 08:00$ – $15:00$ MLT/ 70 – 80° MLAT sector: (i) near-simultaneous activations of BPS auroral forms in the pre- and postnoon sectors, each of which is moving noonward, followed by (ii), PMAFs which are moving northwest or northeast, depending on the B_y polarity, and (iii), a fading phase, characterized by predominantly poleward motion, within the regime of mantle precipitation.

2 Data presentation

2.1 Interplanetary observations

Figure 1 shows interplanetary parameters from spacecraft Wind for the interval 07:00–09:00 UT on 8 December 1999. We have included the ratio $|B_y|/|B_z|$, which gives extra information, in addition to the clock angle, and which past work has shown to be a useful parameter in discussions of ionospheric convection (Knipp et al., 1993; Freeman et al., 1993; Greenwald et al., 1995; Ruohoniemi and Greenwald, 1996).

The 2-h-long interval may be divided into subintervals, ordered by the clock angle, θ , as indicated by the vertical guidelines. We note the following features:

(i) a high speed solar wind ($V \sim 550$ km/s) of stable plasma parameters and dynamic pressure ~ 1.5 – 2 nPa, and (ii) a fairly steady IMF amplitude of strength ~ 5 nT, whose orientation is strongly variable.

The vertical guidelines A–G indicate times we shall examine in detail. In particular, we want to focus on these IMF features:

(i) A–B: strongly northward IMF orientation, (ii) B: sharp southward turning, (iii) C–D: strongly southward IMF ($|B_y|/|B_z| < 1$), (iv) D: B_y and B_z polarity reversals, (v) D–E: B_y negative interval, contrasting with the otherwise predominantly positive B_y , (vi) F–G: excursion from a brief entry into the “intermediate” clock angle regime (90° – 135°) to strongly south ($\theta > 135^\circ$) and back to “intermediate”.

With Wind located behind the dawn terminator at $(-37, -59, 2)R_E$, we estimate a negligible propagation Wind-magnetopause. This estimate is in good agreement with the ground-based observations presented later in the paper. While Wind’s Y-coordinate is large, we note that this is of the order of the estimated coherence length of the solar wind/IMF structures in the Y-direction (Matsui et al., 2002).

2.2 Auroral observations

In this section we report auroral observations corresponding to the IMF data shown in Fig. 1, representing the interval 07:00–09:00 UT on 8 December 1999. Special emphasis will

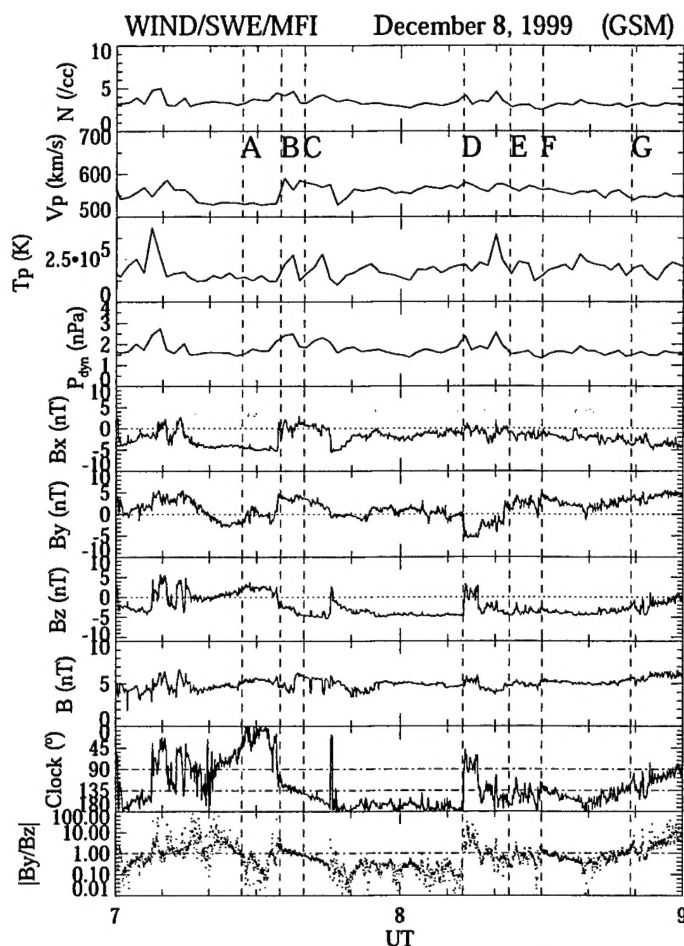


Fig. 1. Solar wind plasma and IMF data obtained from Wind during the interval 07:00–09:00 UT on 8 December 1999. Panels from top to bottom show: proton density, bulk velocity, proton temperature, solar wind dynamic pressure, interplanetary magnetic field (IMF) components B_x , B_y , B_z , the field magnitude, the IMF clock angle, and the ratio $|B_y/B_z|$. Field orientation regimes are marked by vertical guidelines A–G.

be placed on the response to the IMF variability, which is illustrated by the vertical guidelines A–G. The presentation below is mostly, but not always, in chronological order. The exception is the presentation of ASC data for the subinterval of strongly northward IMF orientation (A–B). Since the latter data are included as a contrast to the data representing southward IMF conditions (the main focus in this paper), they come after the presentation of the southward IMF aurora in interval B–C. While meridian scanning photometer (MSP) data are shown in two wavelength channels (630.0 and 557.7 nm) for the whole interval, all-sky camera (ASC) data are shown for more limited intervals, which are selected with the purpose of illustrating main features of the 2-D auroral response. We also note that in the ASC data presentation we apply the following three options: (i) red line only, (ii) green line only, and (iii) both. The reasons are the following. When the features we want to expose are clearly expressed in one channel (the second channel does not add much) we select the channel showing the features most clearly. For exam-

ple, as the BPS and type 2 aurorae appear most clearly in the green line emission, we show that channel only. (Also note that MSP data for the whole period are shown in both emission lines). On the other hand, since most type 1 features, like the PMAFs, typically appear strongly in both emissions, we decided to give one example in both channels.

2.2.1 Interval 07:00–08:10 UT

The most relevant meridian scanning photometer (MSP) and all-sky camera (ASC) data for this interval are shown in Figs. 2 and 3, respectively. The observation geometry and a summary of main features of the observations are given in Fig. 4.

Figure 2 shows colour-coded plots of MSP data, i.e. line of sight intensity versus zenith angle at 630.0 and 557.7 nm for the interval 07:00–08:10 UT, when the ground station swept through the prenoon from $\sim 10:00$ – $11:10$ MLT. The red (630.0 nm) and green (557.7 nm) auroral lines were se-

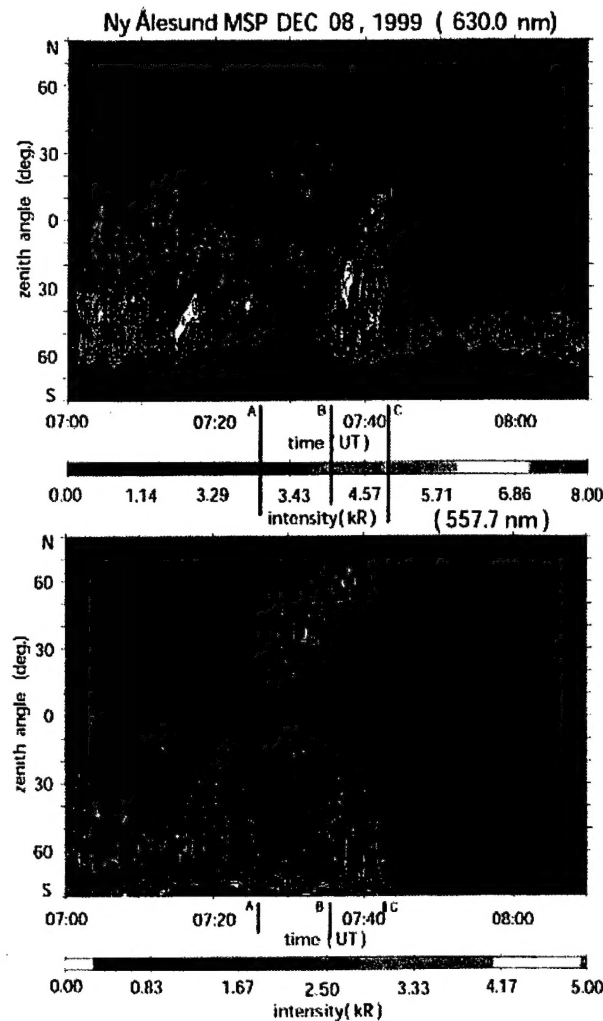


Fig. 2. MSP observations from Ny Ålesund: Auroral line of sight intensities versus zenith angle for the interval 07:00–08:10 UT. The auroral red (630.0 nm) and green (557.7 nm) lines are shown in the upper and lower panels, respectively. Intensity scales in kilo-Rayleighs (kRs) are given at the bottom of each panel. Times corresponding to specific features in the IMF, shown in Fig. 1, have been marked by vertical guidelines (A–C) in the figure. Intensity scale is given at the bottom of each panel.

lected because they are most sensitive to the precipitation of magnetosheath-origin electrons and low-altitude electron acceleration events, respectively (Rees and Luckey, 1974; Rees and Roble, 1986). Auroral events responding to IMF changes marked A, B, and C in Fig. 1 have been indicated in the figure. The top two panels of Fig. 4 summarize the observational phenomenology at the indicated times, the first referring to interval B–C, and the second to the time after C, when the IMF B_z is strongly south.

The strong IMF north orientation during interval A–B is seen to elicit a form to the north of zenith, in both emission lines (Fig. 2). This type 2 aurora is marked as cusp 2 in Fig. 4a.

Before the northward turning at A, during 07:00–07:25 UT, two different auroral types are present. A diffuse, green line aurora is hovering at the southern horizon. This

aurora is absent in the red line emission and is attributed to the precipitation of central plasma sheet (CPS) electrons. It is referred to as CPS in Fig. 4a. Poleward of this is a latitudinally wide region containing two auroral bands, prominent in both emission lines. The poleward boundary of this aurora, which is characterized by strong temporal variability, is located near the zenith of the station (76° MLAT). This aurora is referred to as boundary plasma sheet (BPS) in Fig. 4a. We note that the intensity of this aurora is much reduced in the interval marked A–B, corresponding to strongly northward IMF orientation, when the type 2 cusp aurora is present to the north of zenith.

Next we focus attention on the auroral response to the southward turning of the IMF initiated at B.

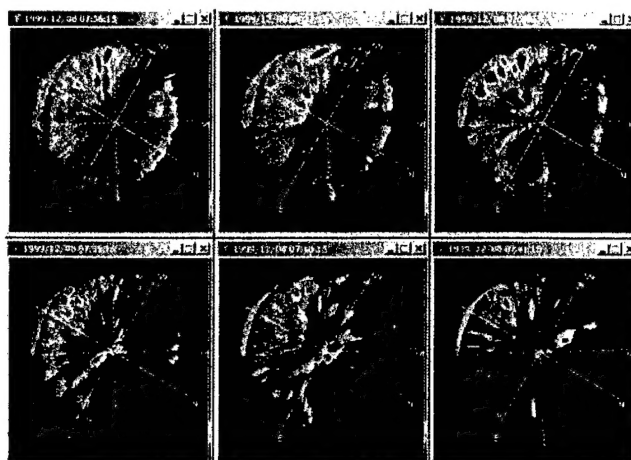


Fig. 3. ASC images at 557.7 nm from Ny Ålesund, representing the times 07:36:15, 07:37:15, 07:38:15, 07:39:15, 07:40:15, and 07:41:15 UT (sequence order: upper panel from left to right, then lower panel in same order). The reference grid is zenith angle (circles at 30, 60, 90°), azimuth angle (straight dashed lines). The MSP scanning meridian (see Fig. 4) is marked north (N) – south (S). The integration time is 1 s.

2.2.2 Southward turning at B–C

The BPS aurora reappeared at full intensity after the southward turning at B (see Fig. 2). Three distinct brightenings are seen in the green line panel of Fig. 2 during interval B–C (07:35–07:42 UT). Figure 3 shows an ASC image sequence of the green line emission covering the interval 07:36–07:41 UT (i.e. B–C, just after the southward turning at B). The white lines indicate the geomagnetic orientation. N–S corresponds to the MSP scanning meridian. The circles give the zenith angle (30, 60, 90°). The field of view (FOV) in the north–south direction is approximately from 70 to 80° MLAT. We note the high degree of fragmentation in the two bands of BPS aurora, located within 30–70° SZ (south of zenith) in the green line emission. The intensification of the southernmost of these bands is accompanied by eastward motion (toward noon). The latter motion is marked by the solid arrow in Fig. 4a.

South of the fragmented bands (BPS), and separated by a precipitation void (particularly clear in the 07:37 UT image), is a more diffuse emission which is of CPS origin, seen clearly both in Fig. 3 and at the southern boundary of the FOV in Fig. 2. The type 2 aurora in the north is evident in Fig. 3 up to 07:39 UT, corresponding to about 3 min after the southward turning of the IMF. Thereafter it is totally absent. The contrast to the strong BPS aurora present in Fig. 3 is seen in Fig. 5, representing interval A–B (IMF strongly north). Here the BPS aurora south of the zenith is very weak, and the type 2 form (northward IMF cusp aurora) to the north (30–50° NZ) is strongly present.

There are three auroral effects at time C, i.e. when the IMF becomes strongly south: (1) the CPS aurora, which had been hovering on the southern boundary, disappears from the FOV. This is synonymous with a shrinkage of the closed field line region, reflecting an earthward displacement of the magnetopause during dayside magnetospheric erosion. (2) The

fragmented BPS bands move out of the MSP FOV to the west (Fig. 4, second panel). (3) Along the MSP meridian these BPS arcs are replaced by a band of red line emission (Fig. 2, top panel). This is a manifestation of the longitudinal expansion of the cusp, as indicated in the Fig. 4b schematic under cusp 1.

2.3 Interval 08:10–09:20 UT

Figure 6 gives MSP data for the interval 08:10–09:20 UT. The red and green line emissions are shown in the upper and lower panels, respectively. We shall focus on the strong auroral brightening marked D, as well as the poleward moving auroral forms (PMAFs) marked F and G. After ~09:05 UT (event H) the aurora configures in two latitudinally separated forms, which is a typical auroral feature during B_y -dominant IMF orientation (not shown in Fig. 1). The latitudinal separation can be seen in both channels during the interval 09:05–09:20 UT. The northernmost branch is rather weak until ~09:15 UT, after which it becomes the dominant branch. The strong activation of the northernmost branch at 09:18 UT (event I) is the response to a sharp northward turning of the IMF, recorded by Wind at 09:20 UT (not shown).

2.3.1 Response to B_y polarity switch at D

The relatively long interval of strongly south IMF orientation (after C) ended by a sharp northward turning at 08:13 UT (Wind time), marked D in Fig. 1. This IMF discontinuity triggered a 5-min long, strong auroral event from 08:14 UT onwards (Fig. 6). After the sharp IMF discontinuity, B_y was strongly negative (~ -5 nT) and B_z slightly positive for ~4 min.

Figure 7 shows ASC red line images for the interval 08:13–08:18 UT. It is seen that the brightening event is activated on the west (prenoon) side of the ~11:30 MLT merid-

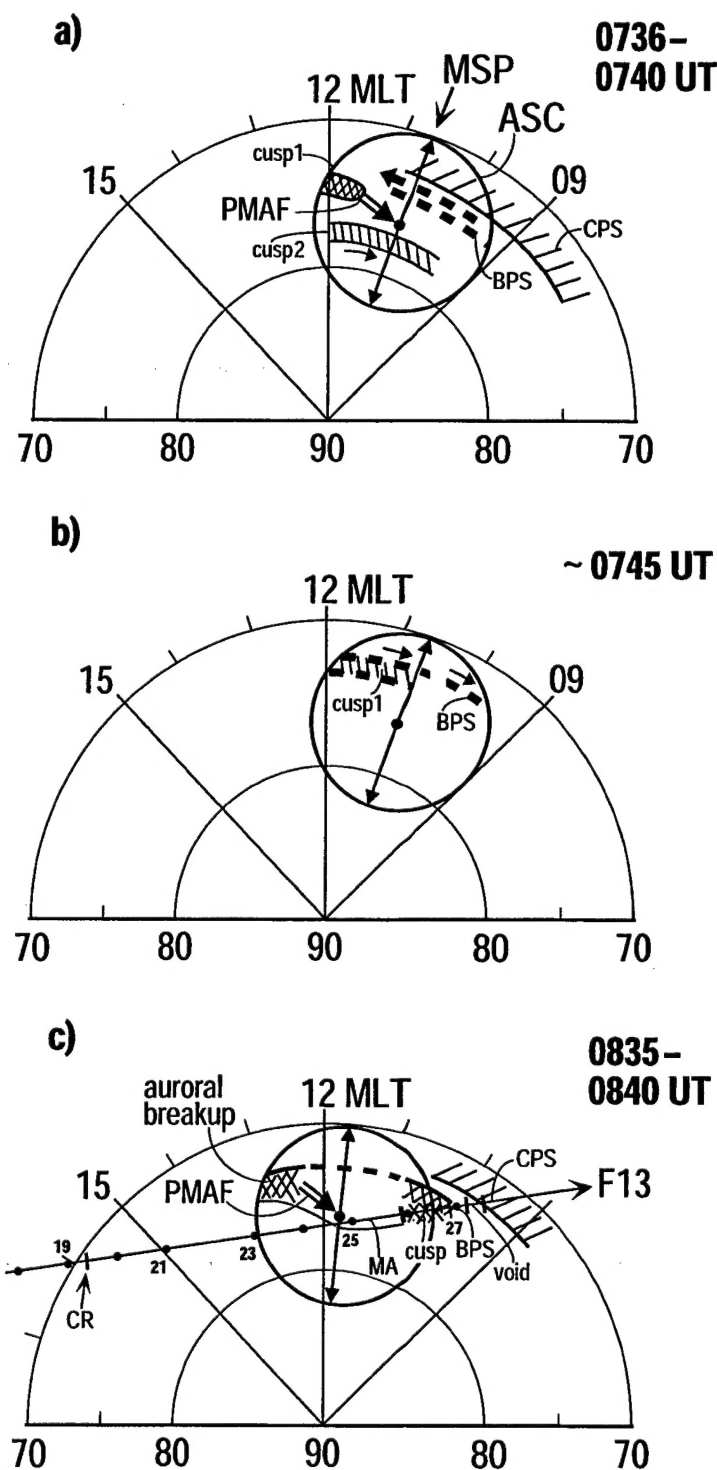


Fig. 4. Schematic overview of auroral observations during the intervals (a) 07:36–07:40 UT (upper panel), (b) ~07:45 UT (middle panel), and (c) 08:35–08:40 UT. The coordinate system is MLT/MLAT. Fields of view of the all-sky camera and the MSP (at 630.0 nm) in Ny Ålesund are marked by the circle and the double-headed meridional line, respectively. The direction of motion of auroral structures is marked by arrows. The track of satellite DMSP F13 during the interval 08:18–08:30 UT is shown in panel (c). A “gap” in the auroral emission in the midday sector is marked by dashed line in the same panel.

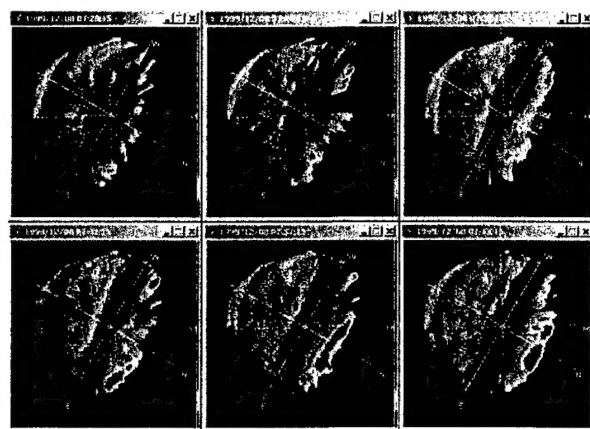


Fig. 5. ASC images at 557.7 nm from Ny Ålesund, representing the times 07:28:15, 07:29:15, 07:30:15, 07:31:15, 07:32:15, and 07:33:45 UT. Same format as in Fig. 3.

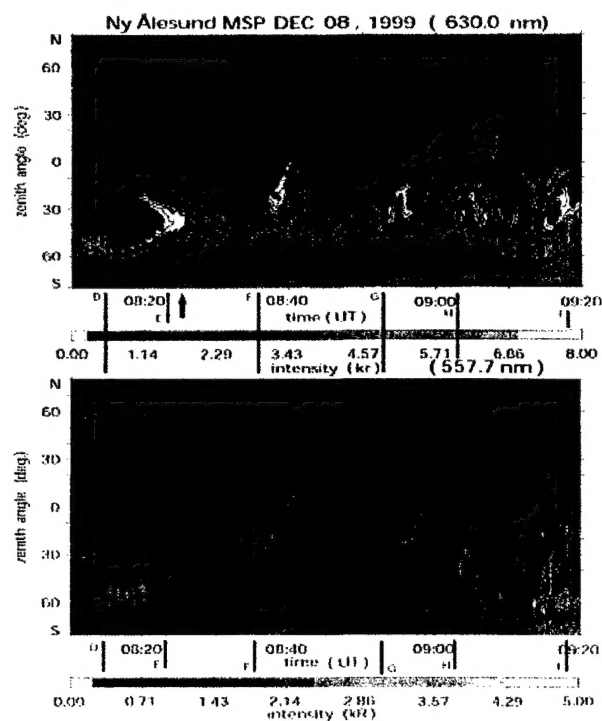


Fig. 6. MSP observations from Ny Ålesund at the wavelength channels 630.0 and 557.7 nm are shown in the upper and lower panels, respectively. Color-coded line of sight intensities versus zenith angle are plotted for the interval 08:10–09:20 UT. Same format as in Fig. 2.

ian (N–S) before expanding rapidly eastward, across noon. This eastward direction of motion of the cusp aurora is attributed to the negative B_y polarity. Examples of westward expanding auroral events during positive B_y conditions are given below.

2.3.2 Events F and G (classical PMAFs)

Figure 6 shows two classical PMAFs marked F and G. The onsets of the two events, as they appear in these MSP records, are separated in time by 17 min. The corresponding IMF θ is $\sim 135^\circ$, and $|B_y|/|B_z| \sim 1$, as shown in Fig. 1.

The red and green line ASC observations for event F are shown in Figs. 8 and 9, respectively. (We note that the colour scales in Fig. 7 and Fig. 9 are different.) As revealed by these illustrations the auroral phenomenon may be divided into two phases. The major initial activation occurs on the eastern boundary of the FOV. A less intense activation is also seen at the western boundary of the FOV. A clear “midday gap” is seen between the prenoon and postnoon activations. This “midday gap” is marked by a dashed line in Fig. 4c. A rapid westward advance of the major activation (postnoon side) occurs during 08:35–08:37 UT, and then a poleward expansion takes place. The two phases may be considered as approximately sequential. As revealed by the green line emission shown in Fig. 9 the event started to fade at 08:40 UT, at the latitude of the ground station (76° MLAT). This expansion of the event is also illustrated in Fig. 4c, specifically the PMAF with an arrow pointing northwest, thus capturing the two components of the motion. We note that the direction of the longitudinal expansion is opposite to that discussed above for the reverse B_y polarity (event D).

Figure 4 shows, in addition, particle precipitation regimes as inferred from a DMSP F13 pass during the interval 08:18–08:30 UT (see below).

2.3.3 Interval between events F and G (08:40–08:53 UT)

We recall from Fig. 1 that this interval corresponds to a strongly south IMF orientation, with $135^\circ < \theta < 170^\circ$ and $|B_y|/|B_z| < 1$. The initiations of the two PMAFs F and G correspond to the boundaries of this interval when this ratio is ≈ 1 and $\theta \sim 135^\circ$.

The auroral state in interval F–G is the same as that for the interval of strongly southward IMF between C and D (see Fig. 2), which was also characterized by an east-west elongated cusp band of moderate intensity (compared to the PMAFs) and the total absence of PMAFs. Thus, this absence is not due to the periodicity of the PMAFs, because its duration (17 min from onset to onset) is much longer than typical PMAF recurrence periods (5–10 min). Rather, it is a response to the external IMF conditions, i.e. the strongly south IMF orientation.

We thus conclude that the upper limit on θ for PMAF activity is around 135 – 150° . For larger clock angles, PMAF activity is replaced by a longitudinally extended band of cusp emission at reduced intensity (Fig. 2,

07:40–08:10 UT; Fig. 6, 08:40–08:53 UT). The auroral response to the strongly south IMF condition corresponds to the “cusp state” which is referred to in the literature as “opening the cusp” (Crooker et al., 1991). The longitudinal expansion in our case is schematically indicated in Fig. 4c. In the same figure (panel c) the reduced auroral intensity during the interval between PMAFs (F–G) is illustrated by a dashed curved line marking the cusp equatorward boundary.

2.4 Ground magnetic disturbance

Figure 10a shows X- and Y-component magnetic deflections from the IMAGE stations on Svalbard (see <http://www.geo.fmi.fi/image/>). The stations are listed according to magnetic latitude, from north to south, spanning the range from 75–71° MLAT. The northernmost station, Ny Ålesund (NAL), is the optical site. The magnetic coordinates (CGM lat.; long.) of the stations are: NAL (75.25; 112.08); LYR (75.12; 113.00); HOR (74.13; 109.59); HOP (73.06; 115.10); BJN (71.45; 108.07).

Events marked in the IMF data (Fig. 1) and showing corresponding signatures in the aurora have been marked by the vertical guidelines A–G in Fig. 10a. We note the following features: (i) the presence of Pc5 magnetic pulsations at the lowest-latitude station (BJN) during 07:00–07:35 UT (up to B), (ii) Pc 5 pulsations at BJN are replaced by a strong convection bay (positive X) at 07:35 UT (B–C), (iii) the high-latitude convection bay (positive X; negative Y) at station NAL ended at 07:40 UT (after southward turning), while the lower-latitude convection bay at BJN–HOR strengthened during the interval 07:35–07:43 UT (B–C), (iv) negative X-deflection beginning (at HOR) at 08:13 UT (D), (v) positive X-deflection at HOR–HOP–BJN at 08:22 UT (E), and (vi) strong convection bays, extending to high latitudes (station NAL), centered at 08:35 UT (F) and at 08:42 UT (G).

We note that the high-latitude event at A–B occurred during intervals of strongly northward IMF orientation, characterized by the presence of a strong type 2 cusp aurora. In a previous study we have shown that this type of magnetic deflection (positive X; negative Y) is due to sunward polar cap (lobe cell) convection (Sandholt et al., 1999).

The two events of positive X-deflection marked F and G are clearly associated with the two major auroral breakup events/PMAFs described in the auroral section. We shall focus on two types of magnetic disturbance. While events F and G are characterized by magnetic deflection of type I (positive X; minor Y), the other type, which we call II (small X; strong, positive Y), occurred during the interval 07:50–08:10 UT (C–D) and around 08:45 UT (between events F and G). The type I and II magnetic deflections correspond to the two specific states/configurations of the cusp aurora, reflecting different IMF orientations, which we refer to as PMAF activity and extended (opened) cusp, respectively. The corresponding convection configurations are (I) the distorted 2-cell, with the presence of strong zonal flow in the midday

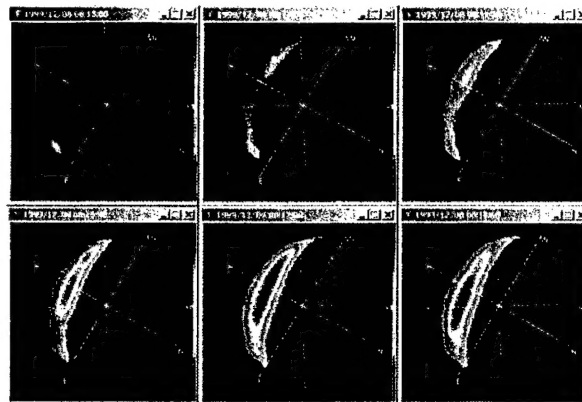


Fig. 7. ASC images at 630.0 nm from Ny Ålesund, representing the times 08:13:00, 08:14:00, 08:15:00, 08:16:00, 08:17:00, and 08:18:00 UT. The reference grid is zenith angle (circles at 30, 60, 90deg) azimuth angle (straight dashed lines). The MSP scanning meridian (see Fig. 4) is marked north (N) — south (S). The integration time is 2 s.

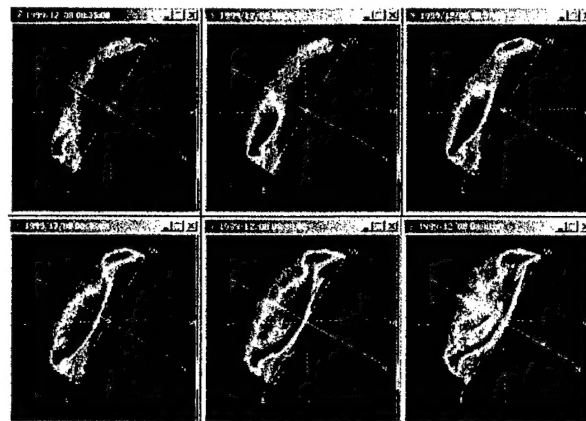


Fig. 8. ASC images at 630.0 nm from Ny Ålesund, representing the times 08:35:00, 08:36:00, 08:37:00, 08:38:00, 08:39:00, and 08:40:00 UT. The format is the same as in Fig. 3.

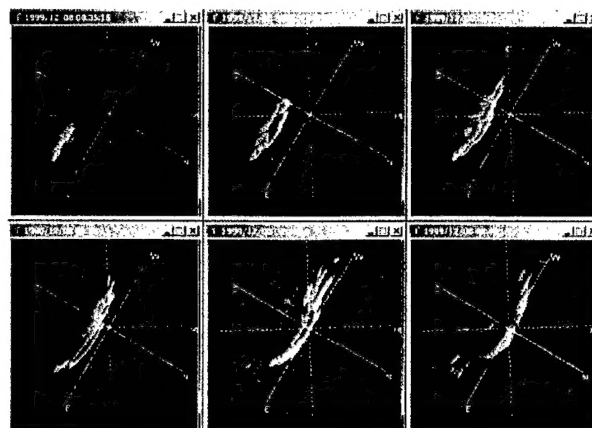


Fig. 9. ASC images at 557.7 nm from Ny Ålesund, representing the times 08:35:15, 08:36:15, 08:37:15, 08:38:15, 08:39:15, and 08:40:15 UT. The format is the same as in Fig. 3.

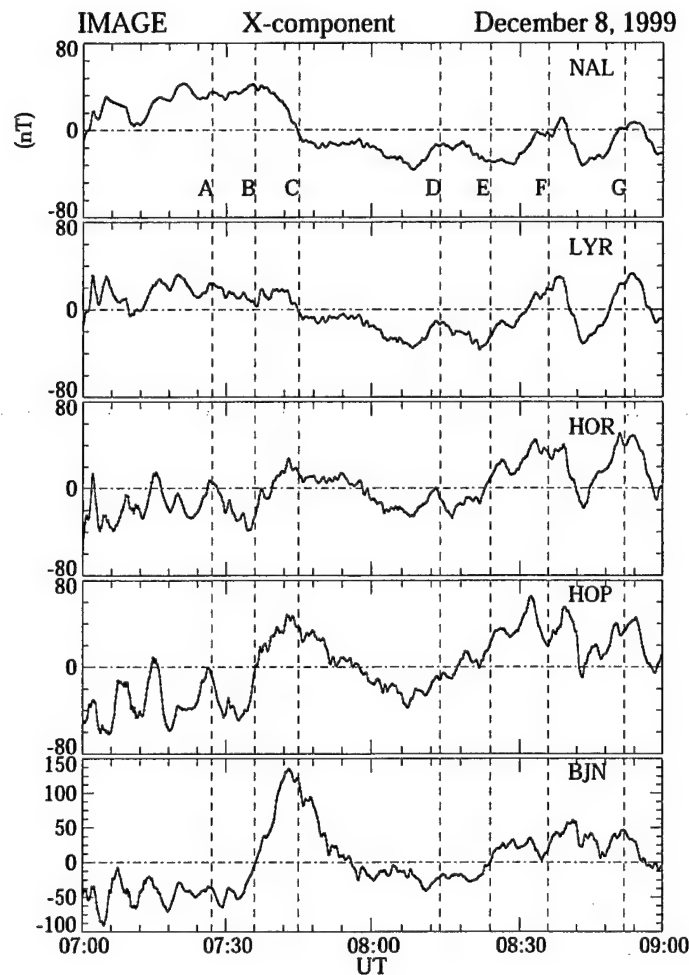


Fig. 10a. Panels (a) and (b) show X- and Y-component magnetic deflections, respectively, at the IMAGE Svalbard stations during the interval 07:00–09:00 UT on 8 December 1999. The stations acronyms are NAL (the optical site), LYR, HOR, HOP, and BJA, spanning the latitude range 76–71° MLAT. The deflection scale is given along the vertical axis on the left side. Events in the IMF data and in the auroral records are marked by vertical guidelines A–G.

sector, and (II) the more symmetrical 2-cell pattern (see discussion section).

2.5 DMSP observations

Figure 11 shows DMSP F13 data obtained during the interval 08:21–08:30 UT on 8 December 1999. The satellite track for the interval 08:18–08:28 UT is indicated in Fig. 4. From the latter figure we can see that the satellite crossed the field of view of the MSP in Ny Ålesund, in the close vicinity of the station, at 08:24:45 UT, near 11:30 MLT/76° MLAT. At this time structured electron precipitation in the energy range 30–400 eV is observed. This is the type of precipitation often observed in the regime called the mantle (poleward of the cusp) during southward IMF conditions (Newell et al., 1991; Newell and Meng, 1994). From the MSP records in Fig. 6 we can see that the cusp aurora at this time (see arrow in the MSP figure) is centered at ~40° south of zenith in Ny

Ålesund (at ~73–74° MLAT). A weak PMAF is approaching the zenith of the station. A patch of cusp-type energy spectrum of the ion-electron plasma is co-located with a channel of enhanced antisunward convection, recorded by the satellite in the 10:30–11:00 MLT/73–74° MLAT sector during 08:25:40–08:26:40 UT. Immediately equatorward of this, we find two narrow dayside BPS arcs (electrons spanning the wide energy range 30 eV–5 keV; average energy ~ 1 keV), then a narrow (~40 km) precipitation void (marked by arrow in the figure), before the entry into the latitudinally wide zone of CPS precipitation, at ~08:27:20 UT. The mentioned precipitation regimes are also marked in Fig. 4. According to Lockwood (1997), the BPS/CPS boundary marks the transition from open to closed field lines. The ~1–10 keV ions in the BPS regime are magnetospheric ions reflected off the interior Alfvén wave (called ri-sp) in the reconnection scenario of Lockwood (1997).

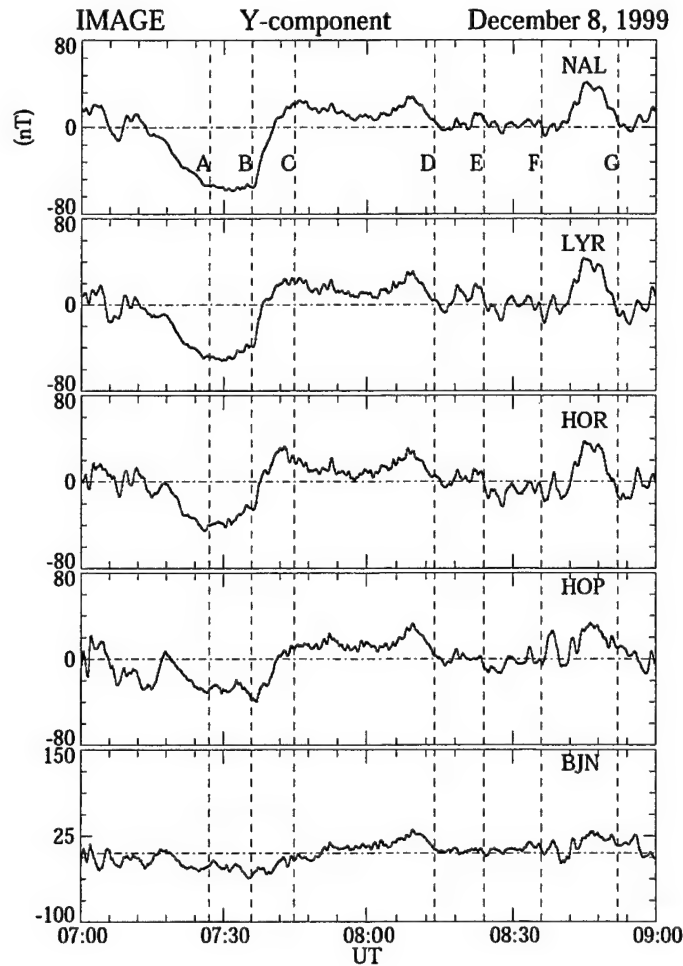


Fig. 10b. See Fig. 10a.

The magnetic deflections recorded from F13 (not shown) indicate the presence of a downward directed cusp current during 08:26–08:27 UT and an upward directed mantle current during 08:24–08:26 UT. The latter current is intensified (steeper gradient) during 08:25:30–08:26:00 UT.

The precipitation recorded from DMSP F13 at $\sim 72.8^\circ$ MLAT/10:00 MLT (08:27 UT), which we refer to as BPS, characterized by keV electrons, corresponds to the aurora (fragmented arcs/bands) called dayside BPS in Fig. 4a, representing 07:36–07:40 UT (see MSP and ASC data in Figs. 2 and 3). We note that the two precipitation features observed at 07:35–07:40 UT and at 08:27 UT, respectively, occur in the same local time sector ($\sim 10:00$ MLT) and correspond to very similar IMF conditions (see Fig. 1).

3 Discussion

3.1 Summary of observations

We have documented in a case study the detailed phenomenology of the dayside auroral response to changing

IMF orientations as given by the clock angle θ and the ratio $|B_y|/|B_z|$. In this study we show examples from the 09:00–12:00 MLT sector. Two distinct auroral configurations are seen corresponding respectively to the (I) “intermediate” clock angle regime (B_y -dominant; θ : 90 – 135°), and (II) the strongly southward-pointing IMF (θ : 135 – 180°).

During the “intermediate” clock angle range, under $B_y > 0$ conditions, the auroral configuration appears as indicated in Fig. 12. Progressing from lower to higher latitudes, we find the following auroral forms:

- (i) At the lowest latitudes ($\sim 70^\circ$ MLAT and below) there is the diffuse aurora of CPS origin.
- (ii) Contiguous to this and on the poleward side is a gap in the green emission, corresponding to a precipitation void (Sandholt et al., 2002a).
- (iii) To the north of this gap appear the fragmented bands (marked by dashed curved lines), corresponding to keV electrons of dayside BPS origin (also referred to as the “soft electron zone”; SEZ), typically moving toward noon, as indicated by arrows. By the high degree of fragmentation, this aurora appears very different from the homogeneous arcs ob-

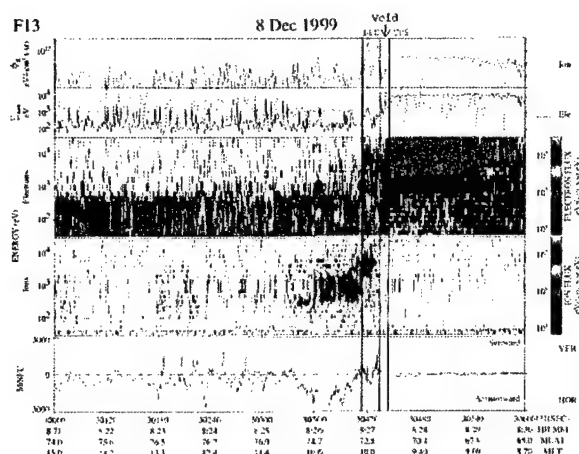


Fig. 11. Particle spectrograms and ion drift obtained from spacecraft DMSP F13 during the interval 08:21–08:30 UT on 8 December 1999. Panels 1 and 2 show differential energy flux for electrons (solid) and ions (dashed) and average energy, respectively. Panels 3 and 4 show color-coded particle spectrograms for electrons and ions, respectively. Panel 5 shows ion drift in the horizontal (cross track; violet) and vertical (upward; green) directions. Different precipitation regimes are marked by vertical guidelines. The precipitation void at the equatorward boundary of the BPS regime is marked by an arrow.

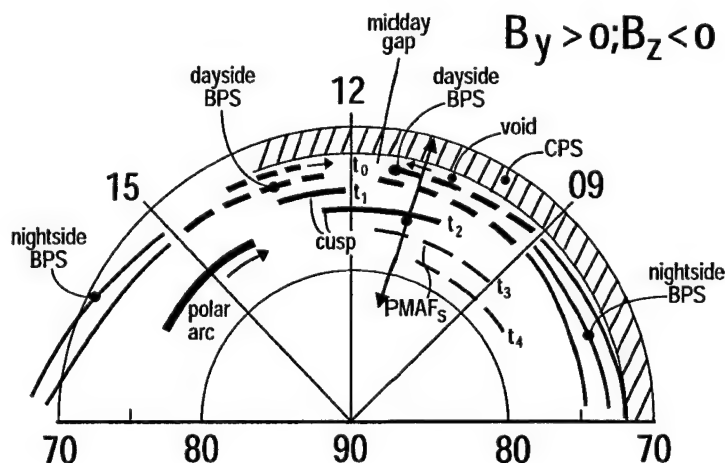


Fig. 12. Schematic overview of the dayside auroral configuration in MLAT/MLT coordinates for the “intermediate” clock angle regime (IMF $B_y > 0$ conditions). Auroral forms are named after their respective particle precipitation sources. Auroral motions are indicated by arrows. The spatio-temporal evolution of the auroral response (PMAF activity) to magnetopause reconnection events (FTEs) is indicated by labels t_0 – t_4 . The field of view of the MSP in Ny Ålesund (at ~08:00 UT) is marked by double-headed meridional line.

served in the morning and evening sectors. The latter forms are marked at nightside BPS in Fig. 12. Our distinction between dayside and nightside BPS aurorae corresponds to the categories corona and multiple arcs, respectively, in the terminology of Yang et al. (2000).

(iv) Poleward of the dayside BPS aurora is the transient PMAF activity (marked cusp 1/PMAFs in the figure). In the case illustrated ($B_y > 0$; Northern Hemisphere) they move westward and poleward. The east-west motion is reversed for the opposite B_y polarity.

A central question in this study is how the dayside auroral pattern responds to southward turnings of the IMF and

the associated reconnection enhancements. In this study we document the following response scenario for different times (t_0 , t_1 , t_2 , t_3 , t_4), when the IMF orientation after the southward turning is characterized by a clock angle in the range 90 – 135° (the “intermediate” state):

t_0 : intensification/activation and noonward motion of the fragmented bands of dayside BPS origin.

t_1 – t_2 : westward and poleward expansion of cusp 1 aurora (PMAFs).

t_3 – t_4 : poleward PMAF motion in the fading stage.

If the IMF after the southward turning is oriented strongly south ($|B_y|/|B_z| < 1$ and $\theta > \sim 150^\circ$), the following au-

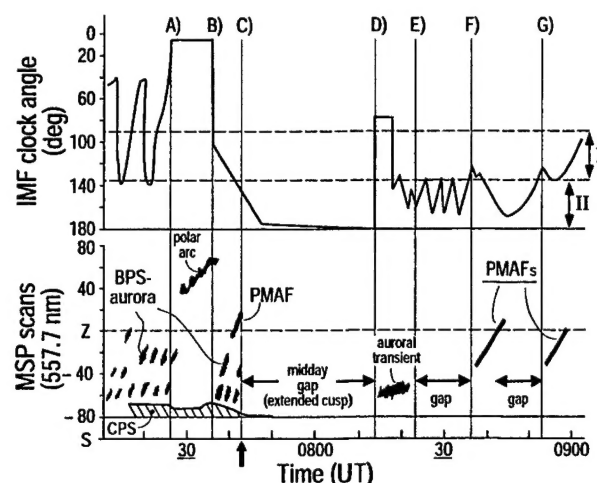


Fig. 13. Schematic overview of the IMF regulation of the green line aurorae, as monitored by the meridian scanning photometer (MSP; see the lower panel). The zenith angle of the MSP records is shown along the vertical axis in the lower panel. North is up. The upper panel shows the essential features of the variability of the IMF clock angle. Clock angle regimes I and II are delimited by horizontal lines. The interval shown is 07:10–09:00 UT, representing the $\sim 10:00$ – $12:00$ MLT sector.

roral characteristics are observed: (i) A weaker and more latitudinally restricted, but longitudinally wide auroral band, compared to the PMAF activity. This state is schematically illustrated in Fig. 4b. The equatorward boundary of this aurora is often located at a lower latitude than the cusp 1 aurora during a B_y -dominant IMF orientation (see, e.g. Sandholt et al., 2003). The more equatorward displacement of the cusp in this state is an indication of strong erosion. This is the aurora corresponding to the expanded cusp (Crooker et al., 1991). Further documentation of the auroral signature of the expanded cusp following a strong southward turning of the IMF, is given by, e.g. Sandholt et al. (1998b).

(ii) The auroral activity referred to as PMAF activity is absent, as exemplified by the interval marked C–D and the 13-min interval between events F and G in the present case. Thus, we find that the PMAF activity is limited to a more narrow clock angle regime (~ 90 – 140°) than usually thought (on this point, see also Drury et al., 2003). By the term PMAF activity we mean an equatorward boundary intensification (EBI), manifest in the BPS aurora, followed by the activation of a cusp form (LLBL/cusp/mantle precipitation) which is moving northwest or northeast, depending on the IMF B_y polarity, as described above. The strong erosion taking place during IMF state II (strongly south) is seen in our case as the equatorward advance of the poleward boundary of the CPS aurora at event C (see lower panel of Fig. 2).

An additional feature to notice in the auroral observations is that the initial appearance of brightening events occur in the prenoon (event D) or postnoon (events F and G) sectors, depending on the IMF B_y polarity (see also Moen et al., 1999). This is consistent with the IMF B_y -related longitudinal shift of the center of the cusp precipitation reported by Newell et al. (1989). These observations may be explained by the partial penetration of IMF B_y into the magnetosphere (see, e.g. Cowley, 1981).

A schematic illustration of the IMF regulation of dayside aurorae during the interval 07:10–09:00 UT, representing the $\sim 10:00$ – $12:00$ MLT sector, is shown in Fig. 13. We note the following features: (i) IMF regulation of the dayside BPS aurora, illustrated by its disappearance during the interval of strongly northward IMF (A–B), and its re-appearance at the time of the rapid southward turning at B, (ii) appearance of a polar arc at the cusp poleward boundary in interval A–B, (iii) equatorward expansion (disappearance from the FOV) of the poleward boundary of the CPS aurora at C (erosion event; marked by arrow in Fig. 13), (iv) the presence of “midday gap” aurora (absence of PMAFs; moderate auroral intensity) during intervals of strongly southward IMF (clock angle regime II), (v) activation of transient type 1 cusp aurora (strong red line emission) at the time of the IMF directional discontinuity at D, (vi) activations of classical PMAFs at the time of transitions of the clock angle from regime II to regime I (at F and G).

3.2 Elements of interpretation

It is suggested that the two dayside auroral configurations present during the “intermediate” clock angle regime (I) and the strongly south IMF orientation (II), correspond, respectively, to two different plasma convection/field-aligned current configurations: (I) a distorted 2-cell versus (II) a 2-cell pattern which is rather symmetrical about noon (see, e.g. Cowley, 2000). This interpretation is supported by the ground magnetometer data reported in this paper. The magnetic disturbance during the intervals of strongly southward IMF (07:50–08:10 UT and 08:40–08:50 UT) are very similar (X component dips; strong, positive Y deflection) and different from that during the “intermediate” clock angle regime (interval B–C and events F and G). The strong, positive Y (eastward) deflection during intervals of strongly southward

IMF orientation is consistent with the expected antisunward convection (sunward Hall current) in a symmetrical 2-cell pattern. Events F and G are, on the other hand, characterized by positive X -deflections, and minor Y -deflection, which reflect a convection pattern (convection currents) containing a strong (dominant) zonal component (westward convection/eastward Hall current).

The rather symmetrical 2-cell convection pattern and longitudinally expanded cusp during strongly southward IMF may imply less intense field-aligned currents (FACs) at noon (see, e.g. Iijima and Potemra, 1976b,a, 1978), compared to the high current intensities, in the form of strong sheet currents, associated with the zonal flow, during the intervals when a strong IMF B_y component is present. The reduced local current intensity near magnetic noon is consistent with the depressed auroral intensity observed in interval C–D, compared to that during the events of PMAF activity (F and G).

Previously, it has been pointed out that the IMF B_y polarity is essential for the motion pattern of PMAFs. The present observations highlight the importance of IMF B_y ($|B_y|/|B_z| \geq 1$) for the occurrence of PMAF activity. According to this interpretation, the 17-min long interval between onsets of events F and G is not due to a natural (intrinsic) recurrence time of PMAFs, which is much shorter (Fasel, 1995), but is rather due to the fact that the required external condition for PMAF occurrence ($|B_y|/|B_z| \geq 1$) is not generally satisfied during this interval. In view of the previously established association between FTEs and PMAFs (Sandholt et al., 1986; Lockwood et al., 1990; Elphic et al., 1990; Fasel, 1995), the observation of the absence of PMAFs during intervals of small or zero IMF B_y component, may indicate that a so-called guidefield along the reconnection line (see, e.g. Siscoe et al., 2001) is a necessary condition for transient reconnection. We note, however, that such an effect has not been seen in statistical studies of the IMF regulation of FTE occurrence (Berchem and Russell, 1984; Kawano and Russell, 1997).

An interesting aspect of the observations is the presence of a longitudinal “gap” (weak background) in the auroral emission, located between clear cusp activations in the prenoon and postnoon sectors, during the time of PMAF activity (see ASC data in Fig. 8 and Fig. 4c for a schematic illustration). This auroral “gap” seems to be consistent with the predicted ionospheric signature of antiparallel reconnection at the dayside magnetopause (Crooker, 1979; Luhmann et al., 1984; Coleman et al., 2001; Rodger et al., 2003). Furthermore, the existence of a restricted IMF clock angle regime (~ 90 – 140°) which is favourable for PMAF occurrence, may also indicate that this auroral phenomenon is a signature of high-latitude, antiparallel reconnection.

On the high-latitude aurora present during the interval of strongly northward IMF (A–B) we note the following. It seems that this aurora, often referred to as type 2 cusp (see Fig. 4a), involves electron precipitation at higher energy than in the type 1 cusp. In the present case this may be inferred from the strong green line intensity. We suggest two possible reasons: (i) different acceleration mechanisms at high

and low magnetopause latitudes (Frey et al., 2003; Østgaard et al., 2003), and (ii) the presence of polar arcs at the cusp poleward boundary during northward IMF orientation (Sandholt et al., 2001, 2002b). The latter is due to current continuity requirements (see, e.g. Reiff et al., 1978) at the poleward boundary of a plasma convection channel in the northward IMF cusp aurora (see illustration in Sandholt et al., 1998a, their Fig. 10).

Next we discuss the question of the dayside BPS aurora in relation to the open-closed field line boundary. According to Kan et al. (1996), the auroral response to reconnection events (FTEs) consists of two parts, (i) “the poleward motion associated with the configurational change and (ii) the noonward motion associated with the convection itself.” The noonward moving auroral forms, which correspond to our BPS aurora, are located on closed field lines, according to Kan. In this paper we document a strong IMF regulation of the dayside BPS aurora. It is observed to be strongly attenuated by the sharp northward turning marked (A), but is clearly re-activated by the subsequent southward turning (B). These fast responses to northward and southward turnings of the IMF, are, in our view, an indication that the dayside BPS is located on open field lines, in agreement with Lockwood et al. (1996) and Lockwood (1997). Our observations are also in agreement with the statement by Lyons et al. (1996) “that large portions of the soft electron zone (SEZ) are on open field lines.”

A second argument which supports the same interpretation is the observation of the green line gap in the aurora, and the associated precipitation void (see Fig. 11) separating the BPS and CPS aurorae in our case. This phenomenon can be explained by the escape of CPS electrons along newly-opened field lines (Scholer et al., 1982; Lockwood, 1997). In an auroral context, the green line gap has been observed previously by Joe Minow (personal communication, 1992), and by Sandholt et al. (2002a) and Sandholt and Farrugia (2002).

Based on these two independent and consistent observations, we conclude that the dayside BPS aurora, appearing in the form of noonward moving fragmented bands in the $\sim 09:00$ – $15:00$ MLT sector, is an important ionospheric manifestation of the reconnection process. This auroral phenomenon may have interesting implications on the reconnection topology and dynamical behaviour of the reconnection X -line (see, e.g. Moore et al., 2002) and the associated boundary layer (Sonnerup and Siebert, 2003) (see their Fig. 8). Details on this are outside the scope of this presentation and will be the topic of a separate paper.

Finally, we note that the observations reported here are mainly from the prenoon sector (09:00–12:00 MLT). For reasons of completeness we think it is appropriate to round off the paper with a comment on the postnoon sector. The postnoon sector morphology is schematically indicated in Fig. 12. We note that the BPS aurora is rather symmetrical about noon (noonward motions on both sides), which is in contrast to the auroral types 1 and 2 (polar arc) that are characterized by a clear IMF B_y -related prenoon-postnoon

asymmetry. The polar arc in the postnoon sector expands into the midday sector in response to northward turning of the IMF. Further details on this will be reported in a follow-up paper. In that paper we will also document the association between the dayside BPS auroral activity and plasma convection (see also Sandholt et al., 2003) and the relationship with magnetopause flux transfer events (FTEs) (see Neudegg et al., 2001). Here we just note that the postnoon sector polar arc, shown in Fig. 12, is that corresponding to the polar cap arc precipitation containing electron energy fluxes extending up to ~ 1 keV, as has been documented by e.g. Lyons et al. (1996) (see their Plate 1). The same ionospheric phenomenon (particle precipitation and field-aligned current structure) has been discussed by Eriksson et al. (2002) as a signature of lobe reconnection.

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